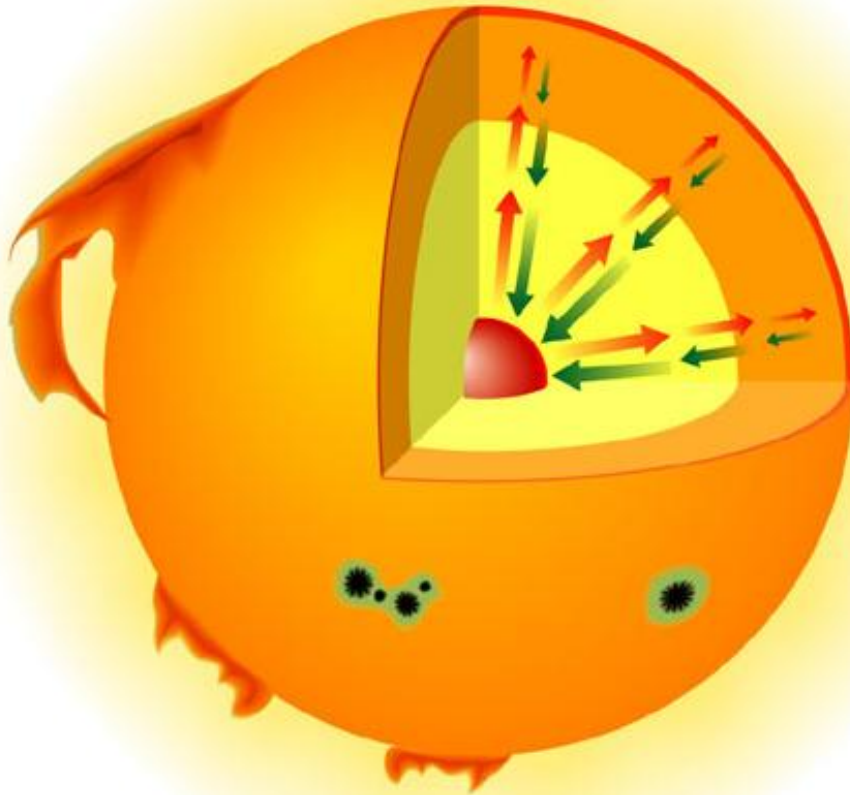


# Star's life: Protracted battle with gravity

pressure →  
gravity ←



**ALWAYS**

To support weight:

⇒ need high pressure

**MOSTLY**

⇒ need high temperature

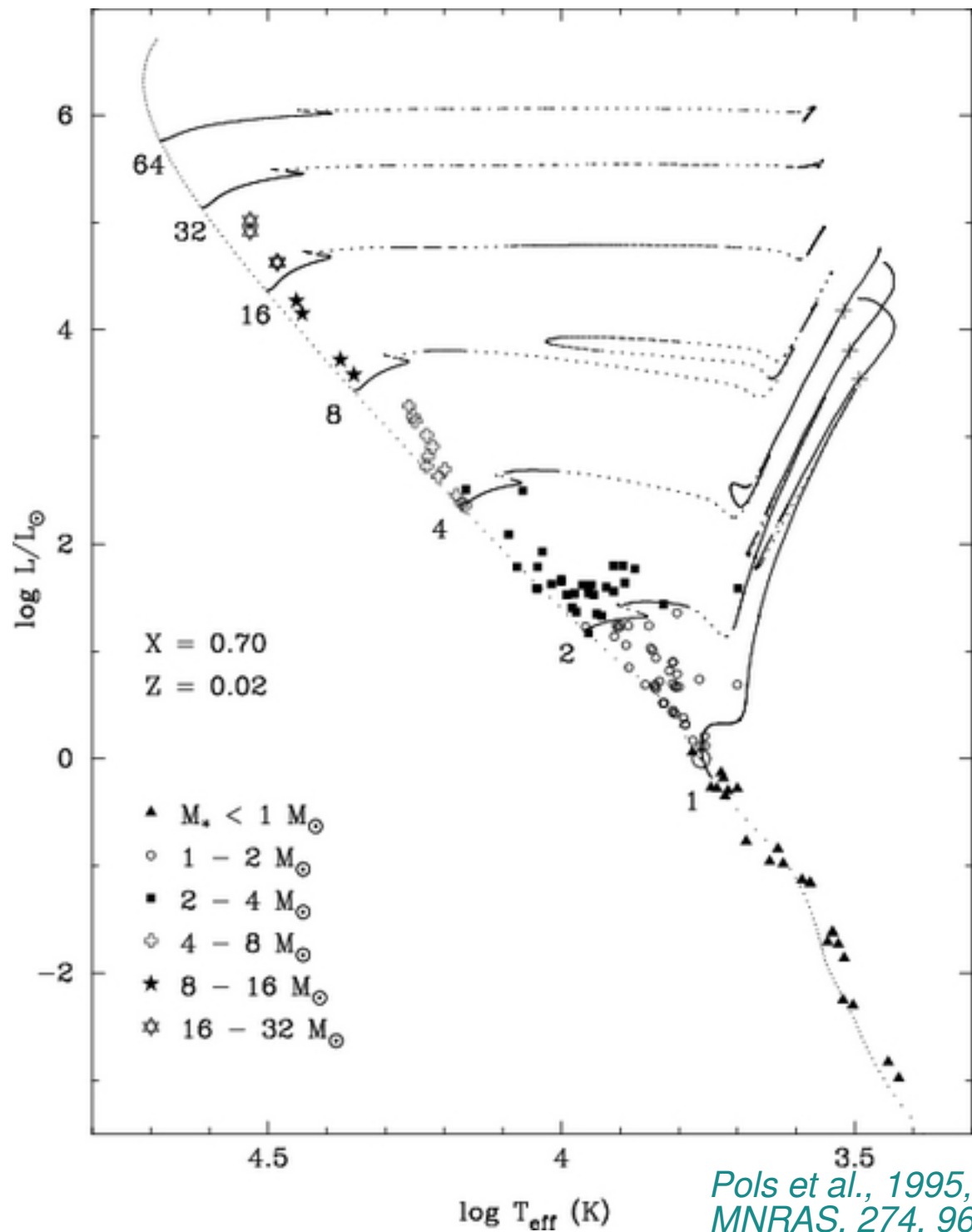
⇒ will lose energy

⇒ need energy source:

- Gravitational contraction
- Nuclear fusion

Ultimately,  
*Can something else than  
thermal pressure balance  
gravity?*

**Evolutionary tracks:**  
what happens depends on mass



*Pols et al., 1995,  
MNRAS, 274, 964*

Distance  $d [\text{pc}] = 1/\pi [\text{arcsec}]$

magnitudes  $m = C - 2.5 \log f$ ;  $m_1 - m_2 = -2.5 \log (f_1/f_2)$ ;  $M - m = 5 - 5 \log d$

Doppler shift  $\frac{\Delta \lambda}{\lambda} = \frac{v_{\text{rad}}}{c}$

Gravity & tides  $a = \frac{GM}{r^2}$ ;  $a_{\text{tide}} \approx \frac{GM}{r^2} \frac{2R}{r}$

Kepler  $GM = \Omega^2 a^3 = \left(\frac{2\pi}{P}\right)^2 a^3$ ;  $\frac{a_1}{a} = \frac{v_1}{v} = \frac{m_2}{M}$ ; for circular orbit,  $v = \sqrt{\frac{GM}{a}}$

Virial Theorem  $E_{\text{kin}} = -\frac{1}{2} E_{\text{pot}}$ ;  $E_{\text{tot}} = E_{\text{kin}} + E_{\text{pot}} = \frac{1}{2} E_{\text{pot}}$  [where  $E_{\text{pot,bin}} = -\frac{GM_1 M_2}{a}$  and  $E_{\text{pot,star}} \approx -\frac{GM^2}{R}$ ]

Ideal gas  $P = nkT = \frac{\rho}{\mu m_H} kT$ ; typical kin. en. per particle  $\langle e \rangle = \frac{3}{2} kT$ ; en. density  $e = \frac{3}{2} nkT = \frac{2}{3} P$

Degenerate gas  $\Delta x \Delta p \sim \hbar$ ;  $E_F = \frac{1}{2} \frac{p_F^2}{m_e} \propto n_e^{2/3}$ ;  $P \propto n_e E_F \propto n_e^{5/3} \propto (\rho/\mu)^{5/3} \rightarrow R \propto M^{-1/3}$

Photon propagation  $l_{\text{mfp}} = 1/\sigma n = 1/\kappa \rho$ ;  $t_{\text{random}} = \frac{R}{l_{\text{mfp}}} \frac{R}{c}$

Black body  $L = 4\pi R^2 \sigma T_{\text{eff}}^4$ ;  $\lambda_{\text{peak}} \propto 1/T$

Hydrostatic eq.  $\frac{dP}{dr} = -\rho \frac{GM}{r^2}$ ;  $\rightarrow P \propto M^2/R^4$

Radiative transfer  $\frac{dT}{dr} = \frac{3\kappa\rho}{16\sigma T^3} \frac{L_r}{4\pi r^2}$ ;  $\rightarrow L \propto T_c^4 R^2 \frac{l_{\text{mfp}}}{R}$

Timescales  $\tau_{\text{ff}} \sim \sqrt{\frac{r^3}{GM}} \sim \sqrt{\frac{1}{G\rho}}$ ;  $\tau_{\text{KH}} \sim \frac{GM^2/R}{L}$

Hydrogen fusion  $E = mc^2$ ;  $p-p$  start with  $p+p \rightarrow D + e^+ + \nu_e$ ;  
CNO catalyst, start with  $^{12}\text{C} + p \rightarrow ^{13}\text{N} + \gamma$

Hydrogen atom  $E_n = -13.6 \text{eV}/n^2$