Stars are divided according to their stage in life.

Lecture 1: Main-sequence stars: H → He in core
What is the main sequence?
How do radius, luminosity & life-time depend on stellar mass?
Hertzsprung-Russell Diagram
Stellar spectral types, UBV photometry

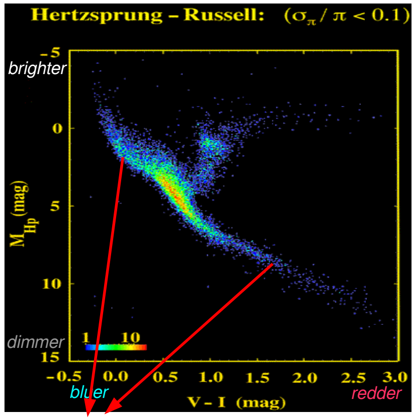
Lecture 2: Giants: high-mass stars (He → C and on to Fe)

Lecture 3: Giants: low-mass stars (H shell, He flash, ...)

Lecture 4: Stellar corpses:
white dwarfs, neutron stars & black holes

Lecture 5: Birth of stars, which also leads to

Lectures 6...: planets...



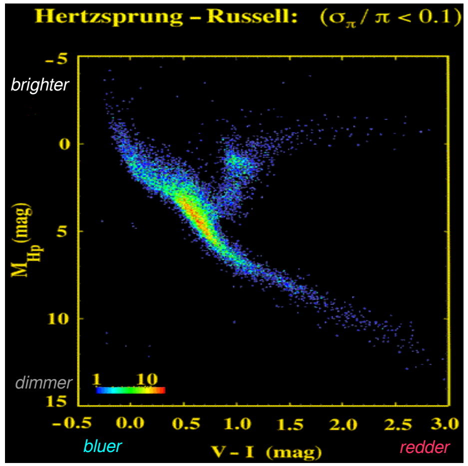
Hipparcos catalogue of nearby stars (with parallax inaccuracy < 10% of parallax measurement) colours indicate density of stars

Main-sequence stars
- Most common
- Stars that are steadily burning H into He in center
- The longest stage in a star's life
- Only small changes in appearance

(most people are adults, fewer infants, adolescents & seniors)

The Color-Magnitude Diagram
(Hertzsprung-Russell Diagram)

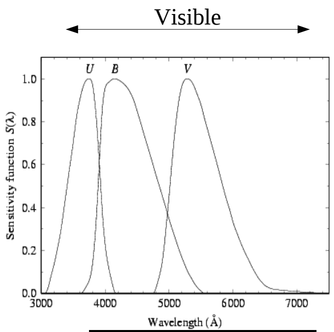
Color is a surrogate for temperature

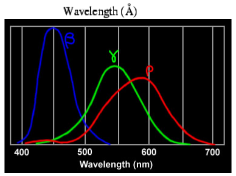


The Color Index

- easier to measure than the entire B.B. Curve
- a good surrogate for T

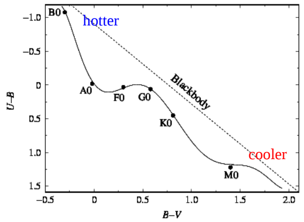
magnitudes defined in photometric bands:
U: ultraviolet
B: blue
V: visual
R: red
I: infrared
(continuing with J, H, K, L, M.....)
Many systems; e.g., in optical, also u,g,r,i,z





Sensitivity for...?

Color-Color plot



Vega (~10,000K) defines U-V = B-V = 0
Sun (~5770K): U-V=+0.16, B-V=+0.64

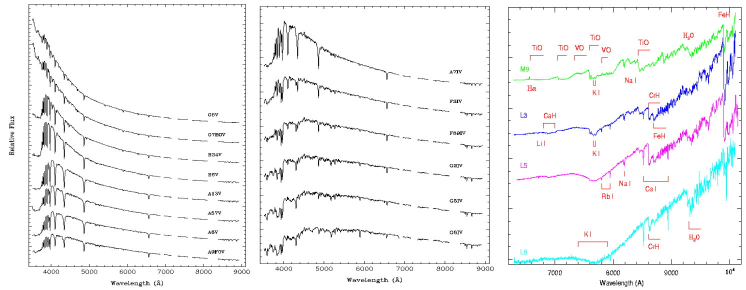
Spectral Types of Stars

One can also use spectral types as surrogates for stellar surface temperature

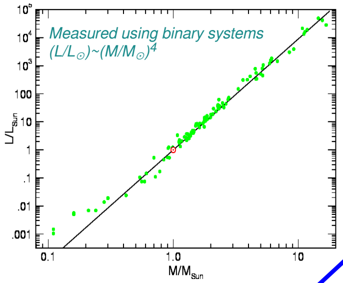
1) historically based on the presence/strength of certain spectral lines (H,He,Ca,Fe...)
2) later shown to be related to temperature

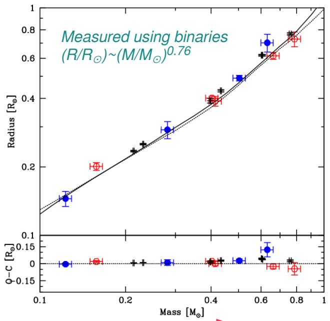
	O	B	A	F	G	K	M	(L	T)
	O0-----O9	B0----B9	A0......A9	F0......F9	G0.....G9	K0.....K7	M0.....M9		
T_{eff}	O5:40,000K		A0:10,000K		G2:5800K		M0:4000K		T:<1000K
	H all ionized				the Sun		molecular absorption bands		

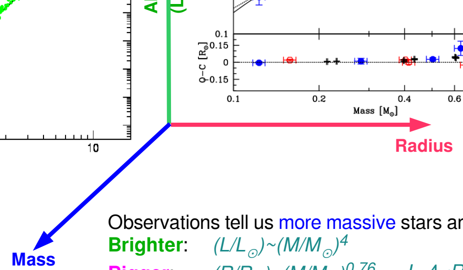
O5-A9 (helium,hydrogen) *A7-G8 (hydrogen, metals)* *M9-L8 (molecular bands)*



Quantifying a main-sequence star
(surface quantities)







Observations tell us more massive stars are
Brighter: $(L/L_\odot)\sim(M/M_\odot)^4$
Bigger: $(R/R_\odot)\sim(M/M_\odot)^{0.76}$ $L=4\pi R^2 \sigma T^4$
Hotter: $(T/T_\odot)\sim(M/M_\odot)^{0.6}$

More massive stars are

Brighter: $(L/L_\odot)\sim(M/M_\odot)^4$
Bigger: $(R/R_\odot)\sim(M/M_\odot)^{0.76}$
Hotter: $(T/T_\odot)\sim(M/M_\odot)^{0.6}$

Stars on Main Sequence follow a 1-parameter sequence

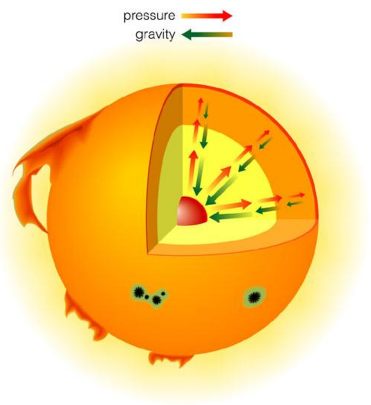
All described using

hydrostatic equilibrium
mass conservation
nuclear burning
photon propagation (opacity)

Stars are modeled numerically:

input parameters:	M_* and chemical composition
equation of state:	$\rho = \rho(P, T, \text{composition})$
energy production:	$\epsilon = \epsilon(\rho, T, \text{composition})$
opacity:	$\kappa = \kappa(\rho, T, \text{composition})$
Solve for:	$P = P(r),\ T=T(r),\ M=M(r),\ L=L(r)$

Star's life: Protracted battle with gravity



ALWAYS

To support weight:

⇒ need high pressure

MOSTLY

⇒ need high temperature

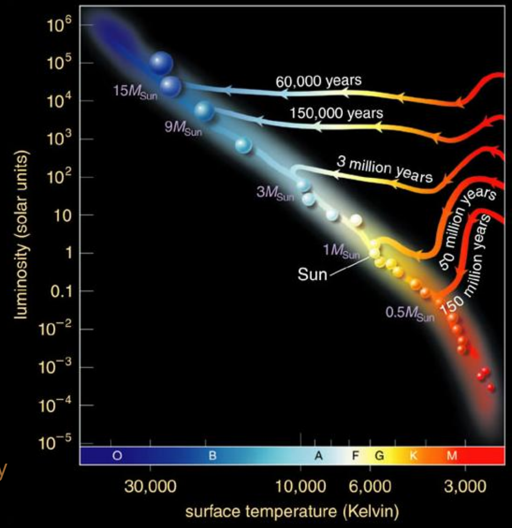
⇒ will loose energy

⇒ need energy source:
- Gravitational contraction
- Nuclear fusion

Ultimately,
Can something else than thermal pressure balance gravity?

Protostellar life tracks

Contraction
⇒ Heating
⇒ Ignition
Or: Ignominy



Why are more massive stars bigger? $(R/R_\odot)\sim(M/M_\odot)^{0.76}$

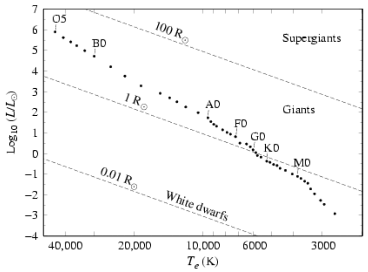
1) Assume all main-sequence stars have the same T_c

$$\text{Virial Theorem: } T_c \propto \frac{M}{R}$$

(massive star has slightly hotter T_c ...)

2) mean density of the star

$$\bar{\rho} \sim \frac{M}{R^3} \propto M^{-2}$$



Low mass limit

For sufficiently small stellar mass (~ 0.08 M_o), electron degeneracy sets in, and the material is no longer as compressible (e.g., a piece of metal)

$$\text{Pressure } P \propto \rho^{5/3} \quad \rightarrow \quad R \propto M^{-1/3}$$

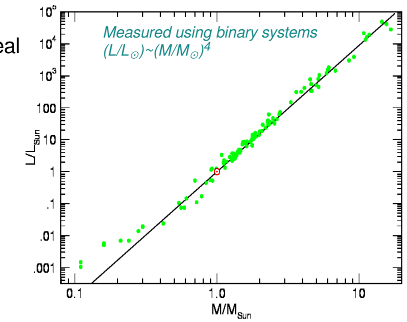
No hydrogen fusion possible ---> Brown Dwarfs

(at about 0.013M_o ~13 M_J, we encounter yet another division: planet/star)

Why are more massive stars brighter? $(L/L_\odot)\sim(M/M_\odot)^4$

Follows from
- hydrostatic equilibrium and ideal gas law (stars are hot)
- equation of radiative transfer (heat leaks out to space)

See PS III.1!
(use constant opacity, e.g., electron scattering, good for hot stars)

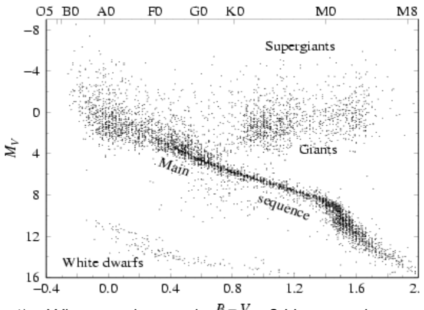


- More massive stars have shorter main-sequence life-time:
  Available nuclear fuel ~ M
  Efficiency of Hydrogen burning 0.7%
  Life-time ~ 0.007*Mc^2/L scales as $1/M^3$

- Very low mass stars are fully 'convective'
  Radiation (slow photon diffusion) replaced by (more efficient) turbulence;
  → luminosity higher than expected based on just photon diffusion

Summary

1) On MS, Mass (or T_{eff} or L) tells all (small variations due to metallicity, age)
2) Hertzsprung-Russell Diagram (HRD, or color-magnitude diagram, CMD)
3) Observational proxies for Luminosity and temperature: M_V, *B-V (V-I,…)*, spectral types



HRD for stars in the Solar neighbourhood

4) What are these other stars? How are they related to each other?
   ------ next lecture